

The miniJPAS survey: stellar atmospheric parameters from 56 optical filters

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ABSTRACT

With a unique set of 54 overlapping narrow-band and two broader filters covering the entire optical range, the incoming *Javalambre-Physics of the Accelerating Universe Astrophysical Survey* (J-PAS) will provide a great opportunity for stellar physics and near-field cosmology. In this work, we use the miniJPAS data in 56 J-PAS filters and 4 complementary SDSS-like filters to explore and prove the potential of the J-PAS filter system in characterizing stars and deriving their atmospheric parameters. We obtain estimates for the effective temperature with a good precision (< 150 K) from spectral energy distribution fitting. We have constructed the metallicity-dependent stellar loci in 59 colours for the miniJPAS FGK dwarf stars, after correcting certain systematic errors in flat-fielding. The very blue colours, including $uJAVA - r$, $J0378 - r$, $J0390 - r$, $uJPAS - r$, show the strongest metallicity dependence, around $0.25 \text{ mag dex}^{-1}$. The sensitivities decrease to about 0.1 mag dex^{-1} for the $J0400 - r$, $J0410 - r$, and $J0420 - r$ colours. The locus fitting residuals show peaks at the $J0390$, $J0430$, $J0510$, and $J0520$ filters, suggesting that individual elemental abundances such as [Ca/Fe], [C/Fe], and [Mg/Fe] can also be determined from the J-PAS photometry. Via stellar loci, we have achieved a typical metallicity precision of 0.1 dex. The miniJPAS filters also demonstrate strong potential in discriminating dwarfs and giants, particularly the $J0520$ and $J0510$ filters. Our results demonstrate the power of the J-PAS filter system in stellar parameter determinations and the huge potential of the coming J-PAS survey in stellar and Galactic studies.

Key words: methods: statistical – techniques: photometric – stars: abundances – stars: fundamental parameters.

1 INTRODUCTION

The *Javalambre-Physics of the Accelerating Universe Astrophysical Survey* (J-PAS; Benítez et al. 2014) aims to image 8500 deg^2 of the northern sky with a unique set of 54 overlapping narrow-band filters and two broad-band filters, with a dedicated 2.5-m telescope, JST/T250, at the *Observatorio Astrofísico de Javalambre* (OAJ; Cenarro et al. 2014). The narrow-band filters have a typical FWHM of 145 Å , covering the entire optical range from 3785 to 9100 Å . Driven by accurate redshift estimates of galaxies over a wide range of epochs, the J-PAS survey will effectively deliver a low-resolution spectra for every single pixel of the 8500 deg^2 sky area observed, providing valuable data sets for not only cosmology and galaxy evolution, but also stellar physics and near-field cosmology.

As a pathfinder for J-PAS, the miniJPAS survey (Bonoli et al. 2021) has imaged the famous Extended Groth Strip (EGS) field of

about 1 deg^2 , with the JPAS-Pathfinder (JPF) camera. In addition to test the performance of the telescope, the miniJPAS data offers a great opportunity to prove the potential of the J-PAS filter system. In this work, we explore the power of these filters in characterizing stars and deriving their fundamental parameters via different techniques.

Metallicity-dependent stellar locus (Yuan et al. 2015a) provides a simple yet non-trivial tool to investigate the dependencies of colours on metallicity and to determine the metallicity for millions of stars from multi-band photometry. Yuan et al. (2015b) applied this technique to the re-calibrated Sloan Digital Sky Survey (SDSS; York et al. 2000) SDSS/Stripe 82 (Yuan et al. 2015c) and obtained metallicities for half a million FGK dwarf stars, with a typical error of $\sigma[\text{Fe}/\text{H}] \sim 0.1\text{--}0.2 \text{ dex}$. Using the same data, Zhang et al. (2021) obtained metallicity-dependent stellar loci for red-giant stars and used them to derive metallicities of giants to a precision of $\sigma[\text{Fe}/\text{H}] \sim 0.20\text{--}0.25 \text{ dex}$, and to discriminate metal-poor red giants from main-sequence stars based on SDSS photometry. With the recalibrated SkyMapper Southern Survey (SMSS) DR2 (Huang et al. 2021), Huang et al. (2022) determined metallicities for over 24-million stars with a similar technique, achieving a precision comparable to

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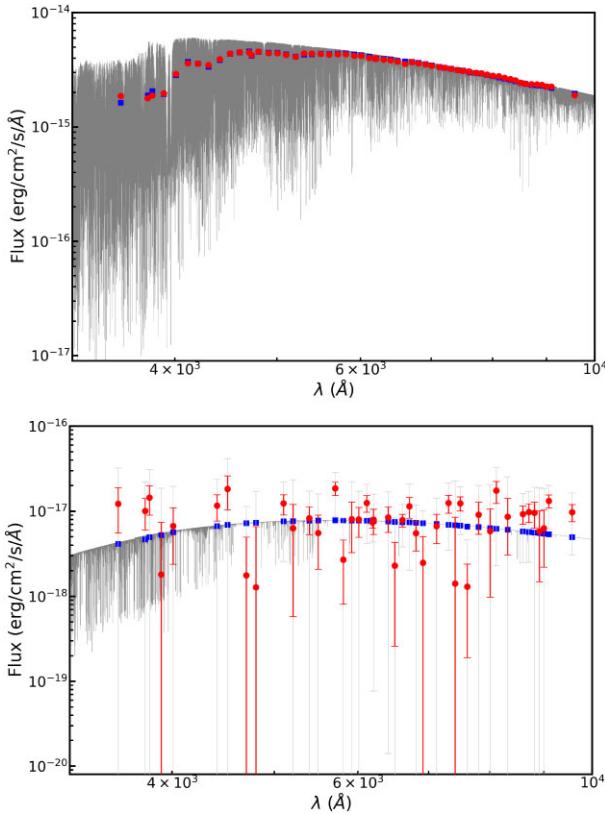


Figure 1. SED fittings obtained with VOSA. The observed J-PAS photometric data are shown in red. The synthetic photometry obtained from the best-fitting BT-Settl model is presented as blue squares, and the corresponding spectra represented by a solid grey line. The light-grey vertical bars are the 3σ uncertainties. Top panel: SED best-fitting model for the 2406-6604 object, with V_{gfb} of ~ 0.07 . Bottom panel: SED best-fitting model for the 2243-8717 object as an example of a rejected source, with V_{gfb} of ~ 3.02 , just above the imposed threshold (see text).

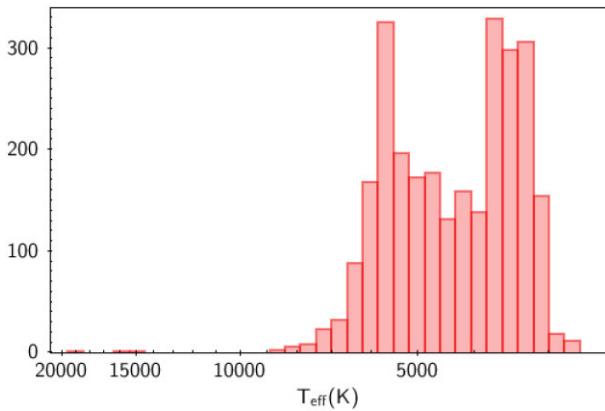


Figure 2. Histograms showing the distribution of T_{eff} obtained from SED fitting with VOSA for the sample with a good fit ($V_{\text{gfb}} < 3$).

or slightly better than that derived from spectroscopy for stars with metallicities as low as $[\text{Fe}/\text{H}] \sim -3.5$. Based on corrected broad-band Gaia Early Data Release 3 (Gaia EDR3; Gaia Collaboration et al. 2016, 2021) colours alone by Niu, Yuan & Liu (2021b) and Yang et al. (2021), and a careful reddening correction, Xu et al. (2022)

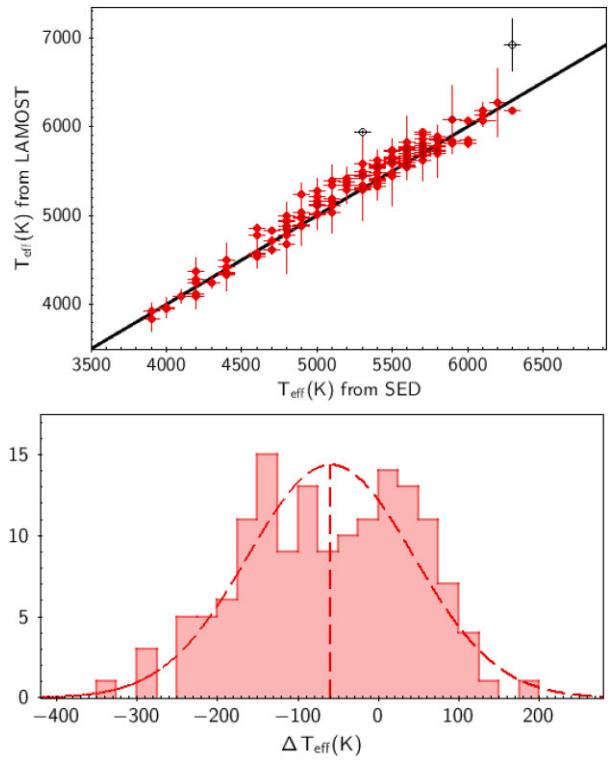


Figure 3. Comparison between effective temperatures obtained from LAMOST spectroscopic data and from miniJPAS SED fitting with VOSA. The black-empty circles are the 3σ clipped outliers. Top panel: Temperatures from LAMOST DR7 (y-axis) and those from VOSA (x-axis), where the solid line is the identity function. Bottom panel: Distribution of T_{eff} differences between VOSA and LAMOST. Dashed line corresponds to a Gaussian fit of the distribution, with a mean T_{eff} value of -60 K and a $\sigma \sim 103$ K, approximately.

determined reliable metallicity estimates for a magnitude-limited sample of about 27-million stars down to $[\text{Fe}/\text{H}] = -2.5$.

In this work, we first determine effective temperatures via spectral energy distribution (SED) fitting, based on miniJPAS photometry. We then investigate the dependencies of miniJPAS colours on the metallicity and perform $[\text{Fe}/\text{H}]$ estimates for FGK dwarf stars via the metallicity-dependent stellar loci. At last, we explore giant/dwarf classifications using a machine learning technique to identify the most sensitive colours. Note that measurements of surface gravity are not discussed in this work.

This paper is organized as follows: Section 2 describes the miniJPAS data used in this work. Section 3 presents the stellar characterization using the available photometric data. The metallicity-dependent stellar loci are determined and discussed in Section 4. Metallicity estimates are described and compared in Sections 5 and 6 discusses giant/dwarf classifications. Finally, we summarize our work in Section 7.

2 MINIJPAS DATA

The miniJPAS was carried out at OAJ with the JPF camera mounted on the JST/T250 telescope. JPF was the first scientific instrument of the JST/T250, before the arrival of the Javalambre Panoramic Camera (JP Cam; Taylor et al. 2014; Marín-Franch et al. 2017). The JPF has a single 9200×9200 CCD, located at the centre of the

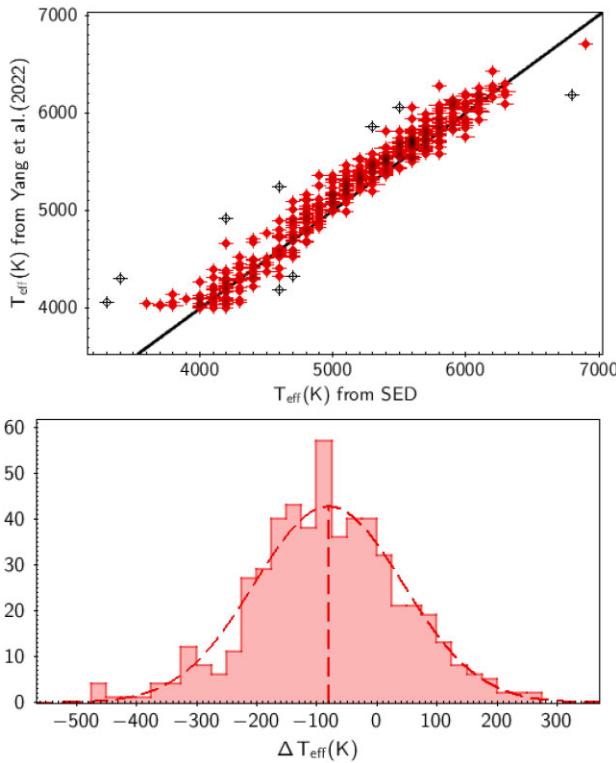


Figure 4. Comparison between effective temperatures from VOSA (this work) to those from Yang et al. (2022) for the common subset, with 542 stars. The black-empty circles are the 3σ clipped outliers. Top panel: Temperatures from previous work (Yang et al. 2022) (y-axis) and those from VOSA (x-axis), where the solid line is the identity function. Bottom panel: Distribution of T_{eff} differences between VOSA and Yang et al. (2022). Dashed line corresponds to a Gaussian fit of the distribution, with a mean T_{eff} value of -81 K and a $\sigma \sim 124$ K, approximately.

JST250 field of view (FoV), with a pixel scale of $0.23 \text{ arcsec pixel}^{-1}$, providing an effective FoV of 0.27 deg^2 .

In addition to the 54 narrow-band filters, ranging from 3785 to 9100 \AA , two broader filters, $J0348$ (also known as $uJAVA$, that covers the ultraviolet edge) and $J1007$ (that covers the red wings), and four SDSS-like broad-band filters ($uJPAS$, g , r , and i) are also used. The J-PAS filters have been optimized to deliver a low-resolution ($R \sim 60$) photospectrum of each imaged pixel/target, therefore probing a large number of key stellar absorption features. The technical description and characterization of the J-PAS filters can be found in Marín-Franch et al. (2012). Detailed information of the filters as well as their transmission curves are available at the Spanish Virtual Observatory Filter Profile Service.¹

The miniJPAS covers the EGS area at $(\alpha, \delta) = (215, +53) \text{ deg}$ with four tiles, with a total area of about 1 deg^2 (see Bonoli et al. 2021, for further details). All data collected are processed and calibrated by the Data Processing and Archiving Unit group at CEFCA.² The image depths achieved – 5σ in a 3-arcsec circular aperture – are deeper than 22 mag for filters bluewards of 7500 \AA and about 22 mag for longer wavelengths. The images and catalogues are publicly available on the J-PAS website.³ The 6 arcsec-aperture magnitudes in the ‘dual mode’

¹<http://svo2.cab.inta-csic.es/theory/fps/index.php?mode=browse&gname=OAJ&gname2=JPAS>

²Centro de Estudios de Física del Cosmos de Aragón (Spain).

³http://www.j-pas.org/datarereleases/minijpas_public_data_release_pdr201912

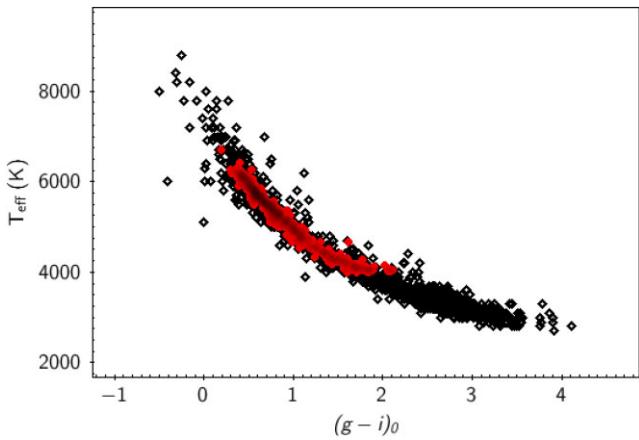


Figure 5. T_{eff} versus $(g - i)_0$ diagram. The black diamonds show the temperatures obtained from VOSA analysis (this work) for 2503 objects – only those objects that presented good g - and i -band photometry (flag < 4 ; $\text{ERR_APER6} < 1$) and VOSA visual good fit ($\text{Vgf}_b < 3$) are shown (see text for details). The red circles represent the temperatures obtained by Yang et al. (2022) for the objects in common (542 stars). Only objects remaining after the 3σ clipping are shown.

catalogues are used in this work, along with aperture corrections. In the ‘dual mode’, the r -band images were used to detect targets with SExtractor (Bertin & Arnouts 1996), and the apertures defined then were also used in other filters.

We select stars from the miniJPAS data base that fulfill the following criteria: $r < 22$ mag and $\text{CLASS_STAR}^4 > 0.6$. A number of 2895 stars are selected, including 1566 and 749 stars brighter than 20 and 18 mag in the r band, respectively. The selected stars are cross-matched with the LAMOST DR7 (Luo et al. 2015) with a match radius of 1 arcsec, 161 common stars are found, including 10 giant stars. The selected stars are also cross-matched with the SDSS/SEGUE catalog (Rockosi et al. 2022), with only 31 stars found. We also cross-match the selected stars with the J-PLUS DR1 stellar parameters catalogue from Yang et al. (2022), 682 stars of $r < 18$ are found. Given the small number of common stars between the miniJPAS catalogue with LAMOST or SDSS spectroscopic catalogues, we have adopted different approaches to characterize the stellar miniJPAS sample, which will be described in the following sections.

3 STELLAR EFFECTIVE TEMPERATURE FROM SEDS

3.1 SED fitting with VOSA

We used the Virtual Observatory Sed Analyzer⁵ (VOSA; Bayo et al. 2008) and the SEDs constructed from the miniJPAS photometry data, to obtain the stellar effective temperatures (T_{eff}) of the objects in the miniJPAS survey via atmospheric model fitting. We used the BT-Settl spectra models by Allard, Homeier & Freytag (2011, and references therein), computed with a cloud model, for all available temperatures – from 400 to $70\,000$ K, with a pace of 100 K till the 7000 K, then a pace of 200 K till the 12 000 K, increasing the pace to 500 K till the 12 000 K, and with a final pace of 1000 K for temperatures larger than

⁴Morphological classification parameter provided by SExtractor.

⁵<http://svo2.cab.inta-csic.es/theory/vosa>

Table 1. Fitting coefficients of flat-fielding errors in 56 J-PAS filters + 4 SDSS-like filters.

Filter	a_0	a_1	a_2	a_3	a_4	a_5
<i>uJAVA</i>	-0.0372	1.65e-06	-4.47e-10	1.20e-05	1.29e-10	-6.57e-10
<i>J0378</i>	-0.0778	8.89e-06	-4.46e-10	1.95e-05	-4.22e-10	-1.22e-09
<i>J0390</i>	-0.0770	1.77e-06	2.38e-10	2.01e-05	-1.05e-10	-9.81e-10
<i>J0400</i>	-0.0391	2.75e-06	-2.59e-10	9.07e-06	2.19e-10	-4.74e-10
<i>J0410</i>	-0.0256	-1.84e-06	2.72e-10	5.84e-06	4.05e-10	-2.76e-10
<i>J0420</i>	-0.0286	4.19e-07	4.51e-11	7.84e-06	2.53e-10	-5.75e-10
<i>J0430</i>	-0.0313	4.15e-06	-2.85e-10	4.71e-06	1.04e-10	-1.16e-10
<i>J0440</i>	-0.0291	2.92e-06	-1.16e-10	5.79e-06	-5.99e-11	-2.30e-10
<i>J0450</i>	-0.0093	-2.30e-06	2.15e-10	2.19e-06	2.81e-10	-3.94e-11
<i>J0460</i>	-0.0137	7.36e-07	-8.81e-11	3.14e-06	9.87e-11	-1.36e-10
<i>J0470</i>	-0.0173	3.78e-06	-4.01e-10	3.67e-06	4.52e-11	-2.91e-10
<i>J0480</i>	-0.0214	4.08e-06	-3.86e-10	2.27e-06	3.70e-11	6.99e-11
<i>J0490</i>	-0.0044	6.19e-07	-4.80e-11	1.12e-06	-6.25e-11	-4.02e-11
<i>J0500</i>	-0.0078	1.74e-07	6.68e-11	1.30e-06	-2.29e-10	1.60e-10
<i>J0510</i>	-0.0214	4.00e-06	-2.89e-10	3.74e-06	-2.37e-10	-6.02e-11
<i>J0520</i>	-0.0166	-3.21e-07	5.52e-11	4.82e-06	-5.34e-11	-1.59e-10
<i>J0530</i>	-0.0021	-6.36e-07	3.83e-11	8.46e-07	-4.93e-11	4.30e-11
<i>J0540</i>	-0.0185	4.97e-06	-4.84e-10	1.39e-06	1.46e-10	-4.10e-11
<i>J0550</i>	-0.0078	2.34e-06	-5.41e-11	-1.22e-06	-1.33e-10	2.84e-10
<i>J0560</i>	0.0020	3.53e-07	-1.38e-10	1.51e-06	-1.74e-10	-1.53e-10
<i>J0570</i>	-0.0001	1.53e-06	-1.64e-10	-2.17e-06	-2.55e-11	3.07e-10
<i>J0580</i>	-0.0158	4.04e-06	-2.89e-10	1.06e-06	-1.67e-10	1.42e-10
<i>J0590</i>	-0.0026	1.22e-06	1.20e-11	-1.41e-06	-1.12e-10	2.29e-10
<i>J0600</i>	0.0033	-7.71e-07	-2.43e-11	-5.27e-07	7.03e-11	6.61e-11
<i>J0610</i>	-0.0125	2.12e-06	-4.45e-11	2.15e-06	-1.14e-10	-1.29e-10
<i>J0620</i>	0.0043	8.08e-07	-9.63e-11	-3.05e-06	-4.09e-11	3.60e-10
<i>J0630</i>	-0.0030	1.24e-06	-5.87e-11	-1.70e-06	-1.39e-11	2.81e-10
<i>J0640</i>	-0.0017	1.12e-06	-1.87e-12	-1.95e-06	-8.76e-11	3.01e-10
<i>J0650</i>	-0.0058	3.17e-06	-2.46e-10	-1.27e-06	-1.04e-10	2.26e-10
<i>J0660</i>	0.0074	-6.19e-07	-4.25e-11	-2.64e-06	7.10e-13	3.29e-10
<i>J0670</i>	0.0035	5.15e-08	-7.59e-11	-1.53e-06	-2.65e-11	2.18e-10
<i>J0680</i>	0.0127	-5.86e-07	7.65e-11	-4.57e-06	-3.76e-11	3.78e-10
<i>J0690</i>	0.0103	-6.42e-07	9.98e-11	-4.43e-06	-2.26e-11	4.29e-10
<i>J0700</i>	0.0053	5.15e-07	-1.28e-10	-1.85e-06	-2.04e-10	3.11e-10
<i>J0710</i>	0.0005	2.02e-06	-1.49e-10	-2.33e-06	-1.30e-10	2.92e-10
<i>J0720</i>	0.0049	8.84e-08	-1.89e-11	-3.45e-06	-3.45e-11	4.66e-10
<i>J0730</i>	-0.0003	3.09e-06	-3.37e-10	-2.97e-06	7.73e-11	2.63e-10
<i>J0740</i>	0.0063	-8.02e-07	4.86e-12	-1.12e-06	8.82e-12	7.74e-11
<i>J0750</i>	0.0046	-1.65e-07	2.14e-11	-1.31e-06	1.47e-11	4.72e-11
<i>J0760</i>	0.0032	9.41e-07	-4.28e-11	-3.13e-06	4.20e-11	2.80e-10
<i>J0770</i>	0.0173	-1.51e-06	-3.35e-10	-2.64e-06	3.54e-12	3.49e-10
<i>J0780</i>	0.0108	2.76e-06	-2.33e-10	-6.05e-06	-3.44e-11	4.23e-10
<i>J0790</i>	-0.0045	1.65e-06	-2.32e-10	4.11e-07	6.03e-12	3.18e-11
<i>J0800</i>	0.0155	-7.40e-08	-2.09e-10	-5.53e-06	1.46e-10	4.69e-10
<i>J0810</i>	-0.0105	2.13e-06	-2.00e-10	4.77e-06	-2.58e-10	-4.24e-10
<i>J0820</i>	-0.0176	6.45e-06	-2.55e-10	2.52e-06	-3.56e-10	-3.64e-10
<i>J0830</i>	0.0028	4.10e-06	-1.38e-10	-6.03e-06	-3.26e-10	6.54e-10
<i>J0840</i>	0.0214	5.00e-07	-6.66e-11	-9.35e-06	-2.31e-10	9.89e-10
<i>J0850</i>	0.0140	4.48e-06	-7.33e-10	-8.61e-06	1.36e-10	7.96e-10
<i>J0860</i>	0.0003	-2.37e-07	1.36e-10	-2.41e-07	-1.17e-10	3.73e-11
<i>J0870</i>	0.0175	-1.52e-06	1.66e-10	-5.13e-06	7.05e-11	2.78e-10
<i>J0880</i>	0.0107	1.44e-06	-1.08e-10	-5.53e-06	3.56e-11	3.90e-10
<i>J0890</i>	-0.0007	2.92e-06	-1.41e-10	-1.06e-06	-3.21e-10	9.02e-11
<i>J0900</i>	0.0170	-2.36e-06	4.74e-10	-7.06e-06	-2.15e-10	7.36e-10
<i>J0910</i>	-0.0010	3.44e-06	-3.03e-10	-1.11e-06	-3.00e-10	1.57e-10
<i>J1007</i>	0.0040	4.35e-06	-4.19e-10	-4.36e-06	-1.90e-10	4.07e-10
<i>uJPAS</i>	-0.0638	8.69e-06	-4.91e-10	1.15e-05	6.44e-11	-6.18e-10
<i>gSDSS</i>	0.0012	-7.54e-07	3.20e-11	-3.00e-07	8.31e-11	4.43e-11
<i>rSDSS</i>	0.0000	0.00e + 00				
<i>iSDSS</i>	0.0012	-7.54e-07	3.20e-11	-3.00e-07	8.31e-11	4.43e-11

Note. $f(x, y) = a_0 + a_1y + a_2y^2 + a_3x + a_4xy + a_5x^2$, where x and y denote CCD positions. $Mag_{\text{new}} = Mag + f(x, y)$.



Figure 6. Flat-fielding errors in units of magnitude in the CCD focal plane for different miniJPAS colours (see text). The errors are derived by a 2D polynomial fitting as a function of CCD (X, Y) positions to the locus residuals. Note: the scales are different for different colours. Note: $r \equiv r_{\text{SDSS}}$.

20 000 K – with $\log g$ varying from 2.0 to 6.0 dex and the metallicity ([M/H]) from –4.0 to +0.5, both with a pace of 0.5. VOSA also allows to work with a range of extinctions. We set this range from

$A_v = 0$ to 0.1, the maximum Milky Way extinction in the region of the sky covered by the miniJPAS (Bonoli et al. 2021). VOSA unreddens the miniJPAS SEDs using 20 values of A_v within this range in steps

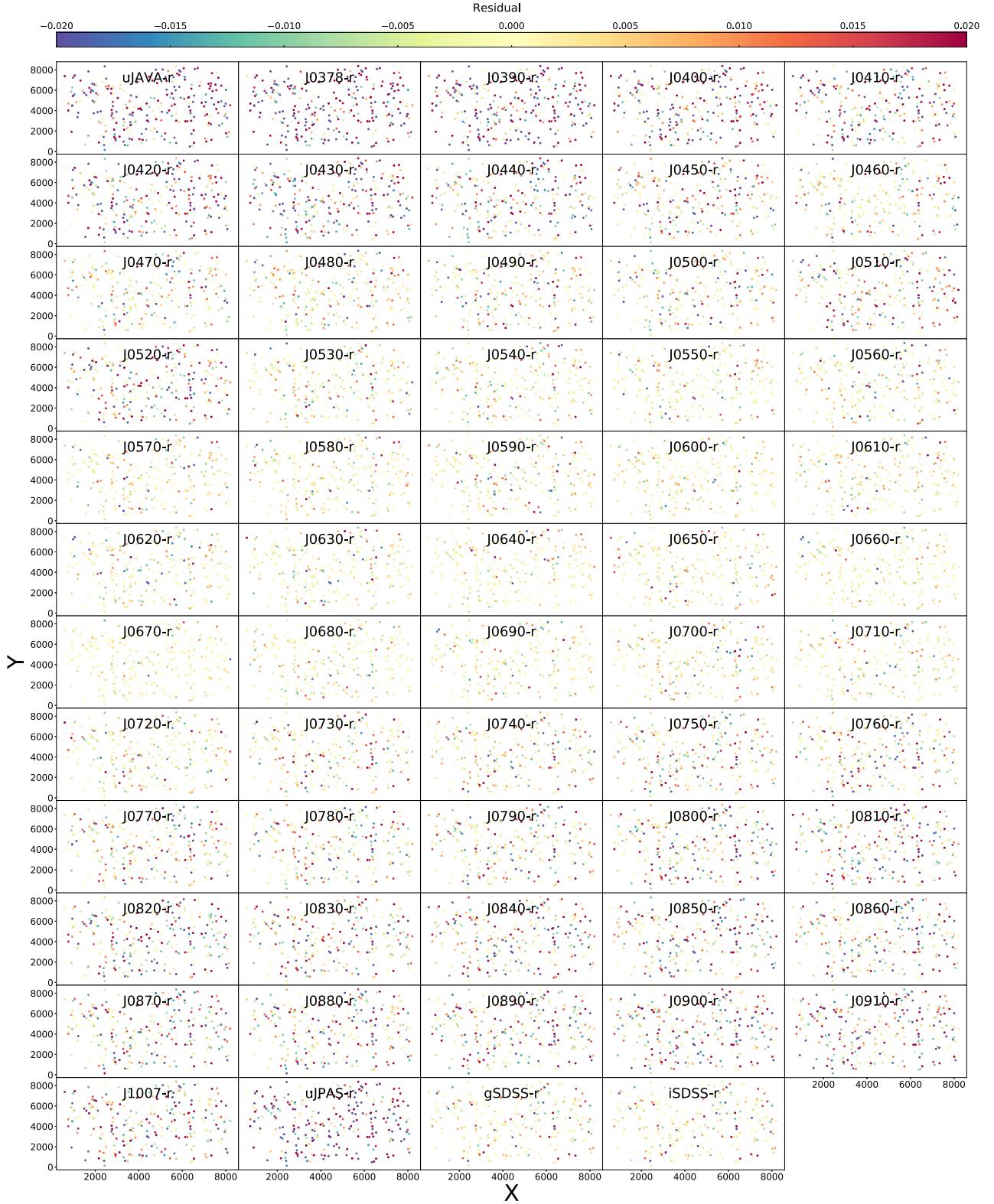


Figure 7. Fitting residuals in the CCD focal plane after correcting for the derived flat-fielding errors. No spatial patterns are found.

of 0.005 mag, and selected the A_v value which unreddened SED best fit the model. It is worth mentioning that fit parameter uncertainties are estimated in VOSA by performing a Monte Carlo simulation with a 100 iterations. Taking the observed SED as the starting

point, VOSA generates 100 virtual SEDs introducing a Gaussian random noise for each point proportional to the observational error. VOSA obtains the best fit for the 100 virtual SEDs with noise and makes the statistics for the distribution of the obtained values for

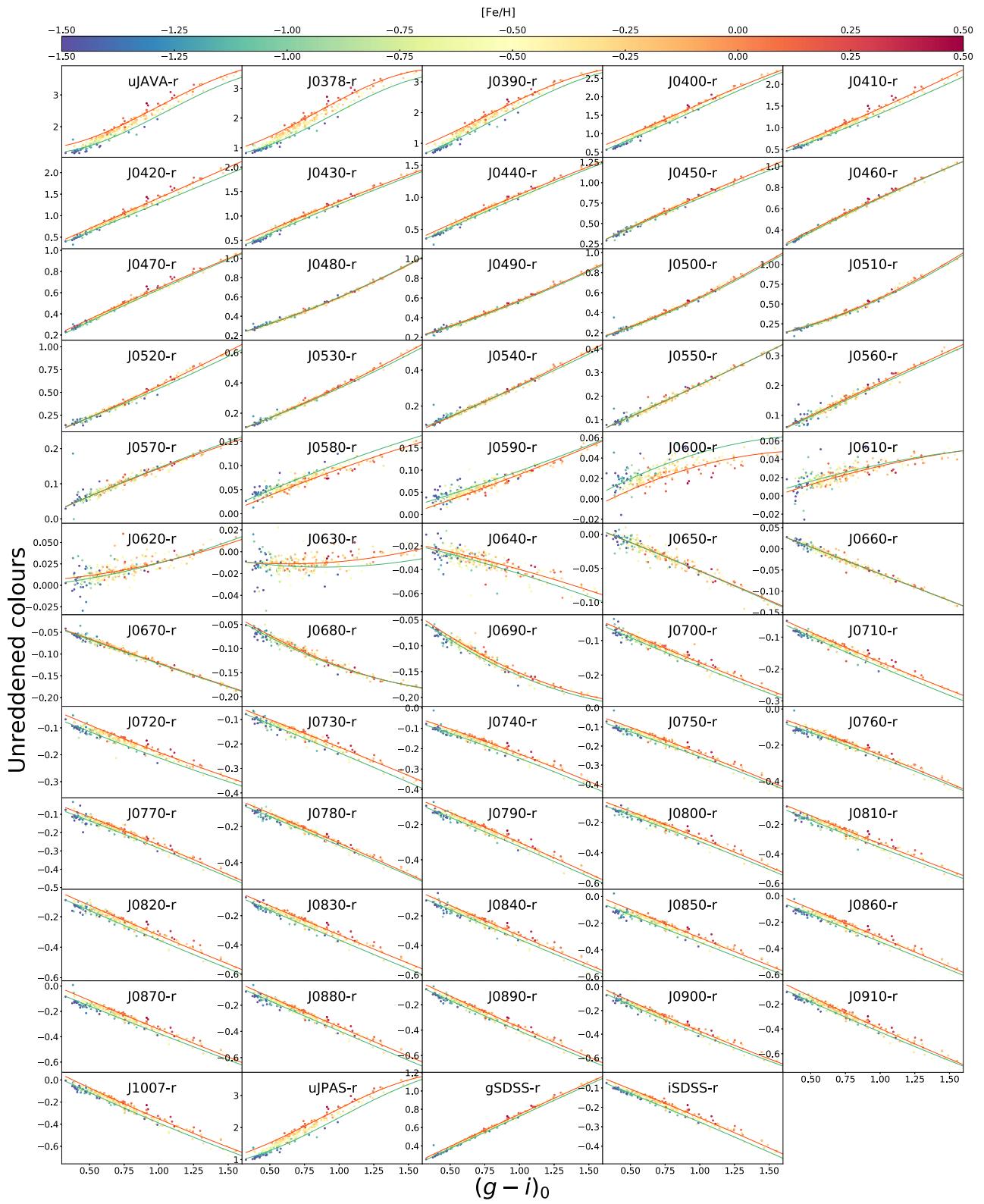


Figure 8. Metallicity-dependent stellar loci of different miniJPAS filters. In each panel, the dwarf stars are colour coded by their metallicities. The red and cyan lines, predicted by the polynomials of Tables 2 and 3, correspond to $[\text{Fe}/\text{H}]$ of 0.0 and -1.0 , respectively.

each parameter. The standard deviation for this distribution will be reported as the uncertainty for the parameter if its value is larger than half the grid step for this parameter.

The complete miniJPAS stellar sample contains 2895 objects. We constructed the SEDs only with good photometric data (individual filter photometry flag < 4). This translated in 2871 miniJPAS SEDs

Table 2. Fitting coefficients of metallicity-dependent stellar loci for FGK dwarf stars in different colours.

Coefficients	<i>uJAVA</i> – <i>r</i>	<i>J0378</i> – <i>r</i>	<i>J0390</i> – <i>r</i>	<i>J0400</i> – <i>r</i>	<i>uJPAS</i> – <i>r</i>
<i>a</i> ₀	1.3497	0.7578	0.4611	0.2539	0.8639
<i>a</i> ₁	0.1453	0.1115	0.4240	0.1342	0.0569
<i>a</i> ₂	0.0542	0.1315	0.1777	–0.0088	0.0107
<i>a</i> ₃	0.0005	0.0155	0.0037	–0.0166	–0.0182
<i>a</i> ₄	–0.6589	0.0724	1.1356	1.2538	0.5197
<i>a</i> ₅	0.3754	0.7846	–0.0411	0.0147	0.6201
<i>a</i> ₆	0.0078	0.0104	–0.1737	–0.0198	0.0094
<i>a</i> ₇	2.9279	2.9310	1.4789	0.4866	1.9751
<i>a</i> ₈	–0.1754	–0.3882	–0.1480	–0.0474	–0.3577
<i>a</i> ₉	–0.9773	–1.1619	–0.6525	–0.1882	–0.7711

Note. $f(x, y) = a_0 + a_1y + a_2y^2 + a_3y^3 + a_4x + a_5xy + a_6xy^2 + a_7x^2 + a_8yx^2 + a_9x^3$, where $x \equiv g - i$ and $y \equiv [\text{Fe}/\text{H}]$.

constructed using between 1 and 60 photometry bands (including the four SDSS-like filters). However, considering the variables on the fitting procedure, a minimum of six photometric points are needed to perform the fit, which reduced the sample to 2858 SEDs, most of them (>94 per cent) with at least 50 photometric points.

From the 2858 analysed SEDs, we found a good-fit solution – defined by a modified reduced χ^2 value (V_{gfb} ⁶) of less than 3 – for 2732 of them, representing ~95 per cent of the miniJPAS analysed sample. To illustrate, Fig. 1 shows two examples of obtained SED fittings, the one with the lowest V_{gfb} , and a rejected solution, close to the cutoff limit of $V_{\text{gfb}} = 3$.

The temperatures obtained range from 2700 to 19 000 K. Fig. 2 shows the T_{eff} distribution obtained from VOSA for the sample with best SED fitting results – 2732 objects with $V_{\text{gfb}} < 3$. The distribution is consistent with what is expected from a field stellar sample, in which dwarfs and giants are present (Van der Swaelmen et al., submitted).

3.2 The LAMOST control sample

In order to evaluate both the accuracy of the fitting method and the reliability of the derived effective temperature, we compare the VOSA results with the effective temperature derived from spectroscopy. We cross-matched the miniJPAS stellar sample with the LAMOST DR7 catalogue (Luo et al. 2015), finding 151 LAMOST objects for which we obtained a good fit with VOSA.

Fig. 3 (top panel) shows the comparison between the T_{eff} obtained by miniJPAS SED fitting with VOSA, and those obtained from spectroscopy by LAMOST. We found a good agreement, within uncertainties, with the SED fitting temperatures slightly cooler than the spectroscopic ones.

Fig. 3 (bottom panel) shows the distribution of the difference in T_{eff} between the values obtained from the miniJPAS SED fitting with VOSA and those from LAMOST spectroscopy, after removing the 3σ clipped outliers. We fitted the distribution to a Gaussian curve, resulting in a mean value of –60 K and a standard deviation of 103 K, approximately. From this, and considering that the LAMOST T_{eff} uncertainties have a median value of ~60 K, we can conservatively estimate a typical T_{eff} uncertainty lower than 150 K in the VOSA miniJPAS SED fitting, in the range of temperatures in common.

⁶ V_{gfb} : Modified reduced χ^2 , calculated by forcing $\sigma(F_{\text{obs}})$ to be larger than $0.1 \times F_{\text{obs}}$, where $\sigma(F_{\text{obs}})$ is the error in the observed flux (F_{obs}). This can be useful if the photometric errors of any of the catalogues used to build the SED are underestimated. V_{gfb} smaller than 10–15 is often perceived as a good fit. <http://svo2.cab.inta-csic.es/theory/vosa/help/star/fit/>.

This result validates the miniJPAS data set and the SED fitting methodology with VOSA as a reliable way to obtain stellar effective temperatures for thousands of sources, without the need of more resource demanding spectroscopic observations.

From the analysed LAMOST sample, we reached the conclusion that even with the amazing collection of 56 narrow-band filters, there is not enough resolution in the SED to derive surface gravities and metallicities with good precision. This result was expected as SEDs have much lower sensitivity to changes in $\log g$ and metallicity. Therefore, other approaches are presented below.

3.3 J-PLUS stellar sample

Given the small numbers of stars in common between the miniJPAS catalogue and the LAMOST/SDSS spectroscopic catalogues, we further used the J-PLUS stellar parameters from Yang et al. (2022) to compare with our miniJPAS SED fitting results. The J-PLUS DR1 catalogue was re-calibrated with the stellar colour regression method (Yuan et al. 2015c; Huang et al. 2021; Niu, Yuan & Liu 2021a; Niu et al. 2021b; Huang & Yuan 2022; Xiao & Yuan 2022), achieving a calibration precision of 2–5 mmag in all 12 J-PLUS filters. Based on the re-calibrated JPLUS DR1 and *Gaia* DR2 (Gaia Collaboration et al. 2018), Yang et al. (2022) trained eight cost-sensitive neural networks, the CSNet, to learn the non-linear mapping from stellar colours to their labels independently. The stellar parameter achieved internal precisions were of 55 K for T_{eff} , about 0.15 dex for $\log g$, and 0.07 dex for $[\text{Fe}/\text{H}]$, respectively. Abundance for $[\alpha/\text{Fe}]$ and 4 individual elements ($[\text{C}/\text{Fe}]$, $[\text{N}/\text{Fe}]$, $[\text{Mg}/\text{Fe}]$, and $[\text{Ca}/\text{Fe}]$) were also estimated, with a precision of 0.04–0.08 dex. The catalogue containing about 2-million stars is publicly available.⁷

We cross-matched J-PLUS DR1 stellar parameters catalogue from Yang et al. (2022) with the miniJPAS sample with good-fitting result from VOSA (2732 objects). A total of 542 stars were found in common. The obtained VOSA temperatures (this work) are in a very good agreement with the previous results by Yang et al. (2022), as illustrated in Fig. 4, where both results are compared. The T_{eff} obtained from the miniJPAS SED fitting with VOSA are presented in the abscissa axis. The temperatures by Yang et al. (2022) are presented in the ordinate axis (top panel), showing the accordance in respect to the identity function (black line). The difference between both results, this work and the results by Yang et al. (2022), are also shown (bottom panel). After applying a 3σ clipping, we found a mean difference of –81 K, with a standard deviation (σ) of 124 K, approximately. These results are in complete agreement, within the

⁷<http://www.j-plus.es/ancillarydata/index>

Table 3. Fitting coefficients of metallicity-dependent stellar loci for FGK dwarf stars in different colours.

colours	a_0	a_1	a_2	a_3	a_4	a_5
$J0410 - r$	0.1153	0.0905	0.0425	1.2298	0.0706	0.0987
$J0420 - r$	-0.0064	0.0735	0.0392	1.4127	0.0837	0.0037
$J0430 - r$	0.0831	0.0990	0.0084	1.2993	-0.0355	-0.0851
$J0440 - r$	0.0802	0.0747	0.0162	1.0258	-0.0198	-0.0621
$J0450 - r$	0.0372	-0.0093	0.0007	0.8531	0.0376	-0.0543
$J0460 - r$	0.0531	0.0315	0.0116	0.7094	-0.0113	-0.0551
$J0470 - r$	0.0374	0.0508	0.0225	0.6404	-0.0122	-0.0336
$J0480 - r$	0.1181	0.0050	0.0044	0.3398	-0.0054	0.1350
$J0490 - r$	0.1000	0.0166	0.0098	0.3847	0.0033	0.0900
$J0500 - r$	0.0339	0.0080	0.0038	0.3861	0.0079	0.1357
$J0510 - r$	0.0180	0.0019	0.0030	0.3172	0.0157	0.2427
$J0520 - r$	-0.0879	-0.0276	-0.0025	0.5926	0.0528	0.0653
$J0530 - r$	-0.0097	-0.0079	0.0000	0.3320	0.0176	0.0521
$J0540 - r$	-0.0265	-0.0144	-0.0008	0.3178	0.0187	0.0238
$J0550 - r$	-0.0184	0.0019	0.0032	0.2551	0.0005	0.0115
$J0560 - r$	-0.0122	0.0068	0.0057	0.2237	0.0046	-0.0029
$J0570 - r$	-0.0223	0.0057	0.0029	0.1810	-0.0052	-0.0154
$J0580 - r$	-0.0234	-0.0101	-0.0008	0.1296	-0.0026	-0.0135
$J0590 - r$	-0.0205	-0.0188	-0.0019	0.1016	0.0091	0.0093
$J0600 - r$	-0.0264	-0.0133	-0.0039	0.0814	-0.0032	-0.0225
$J0610 - r$	-0.0132	-0.0078	-0.0026	0.0566	0.0033	-0.0109
$J0620 - r$	0.0038	0.0077	0.0014	0.0079	-0.0060	0.0145
$J0630 - r$	-0.0049	-0.0046	-0.0019	-0.0193	0.0075	0.0152
$J0640 - r$	-0.0128	0.0014	0.0013	-0.0216	0.0035	-0.0054
$J0650 - r$	0.0281	-0.0014	-0.0004	-0.0789	0.0016	-0.0029
$J0660 - r$	0.0675	0.0022	0.0018	-0.1267	0.0001	0.0002
$J0670 - r$	-0.0072	0.0015	-0.0004	-0.1183	-0.0004	0.0046
$J0680 - r$	0.0152	0.0048	-0.0011	-0.1994	-0.0043	0.0476
$J0690 - r$	0.0142	0.0074	0.0010	-0.2188	-0.0009	0.0521
$J0700 - r$	0.0077	0.0164	0.0047	-0.1811	-0.0014	0.0093
$J0710 - r$	0.0149	0.0167	0.0049	-0.2047	0.0029	0.0115
$J0720 - r$	0.0206	0.0410	0.0123	-0.2286	-0.0083	0.0172
$J0730 - r$	0.0085	0.0197	0.0060	-0.2146	0.0092	-0.0070
$J0740 - r$	0.0071	0.0273	0.0067	-0.2088	0.0004	-0.0243
$J0750 - r$	0.0248	0.0402	0.0109	-0.2444	-0.0117	-0.0187
$J0760 - r$	0.0236	0.0275	0.0024	-0.2567	-0.0097	-0.0220
$J0770 - r$	0.0307	0.0273	0.0030	-0.2818	-0.0068	-0.0165
$J0780 - r$	0.0345	0.0273	0.0047	-0.3148	-0.0096	-0.0112
$J0790 - r$	0.0337	0.0434	0.0115	-0.3238	-0.0087	-0.0115
$J0800 - r$	0.0359	0.0402	0.0118	-0.3174	-0.0062	-0.0225
$J0810 - r$	0.0576	0.0519	0.0178	-0.4069	-0.0082	0.0184
$J0820 - r$	0.0783	0.0477	0.0120	-0.4141	-0.0037	0.0144
$J0830 - r$	0.0717	0.0423	0.0139	-0.4116	0.0066	0.0090
$J0840 - r$	0.0824	0.0453	0.0109	-0.4290	-0.0082	0.0133
$J0850 - r$	0.1161	0.0541	0.0084	-0.4312	-0.0109	0.0061
$J0860 - r$	0.1194	0.0605	0.0100	-0.4315	-0.0177	-0.0046
$J0870 - r$	0.1207	0.0657	0.0100	-0.4755	-0.0226	0.0141
$J0880 - r$	0.0973	0.0354	0.0058	-0.4718	0.0018	0.0095
$J0890 - r$	0.1247	0.0537	0.0131	-0.5238	-0.0112	0.0216
$J0900 - r$	0.1533	0.0524	0.0133	-0.5697	-0.0139	0.0417
$J0910 - r$	0.1343	0.0684	0.0124	-0.5601	-0.0278	0.0288
$J1007 - r$	0.2354	0.0638	0.0173	-0.6406	-0.0097	0.0526
$gSDSS - r$	0.0357	0.0291	0.0069	0.7234	-0.0001	-0.0132
$iSDSS - r$	0.0357	0.0291	0.0069	-0.2766	-0.0001	-0.0132

Note. $f(x, y) = a_0 + a_1y + a_2y^2 + a_3x + a_4xy + a_5x^2$, where $x \equiv g - i$ and $y \equiv [\text{Fe/H}]$.

estimated errors, to those obtained in the comparison to LAMOST data, being VOSA T_{eff} slightly cooler in the 4800–5800 K range. Probably, this trend is due to different systematics between the VOSA and J-PLUS/LAMOST results, since Yang et al. (2022) used LAMOST data for training the artificial neural networks and the same behaviour is seen when comparing the SED T_{eff} obtained

from VOSA to those from LAMOST spectra (see fig. 3, top panel).

The T_{eff} versus $(g - i)_0$ diagram is presented in Fig. 5, where black diamonds show the temperatures obtained from VOSA analysis for the miniJPAS SEDs. Note that only those stars with $T_{\text{eff}} < 10\,000$ K are presented. For comparison, the temperatures obtained

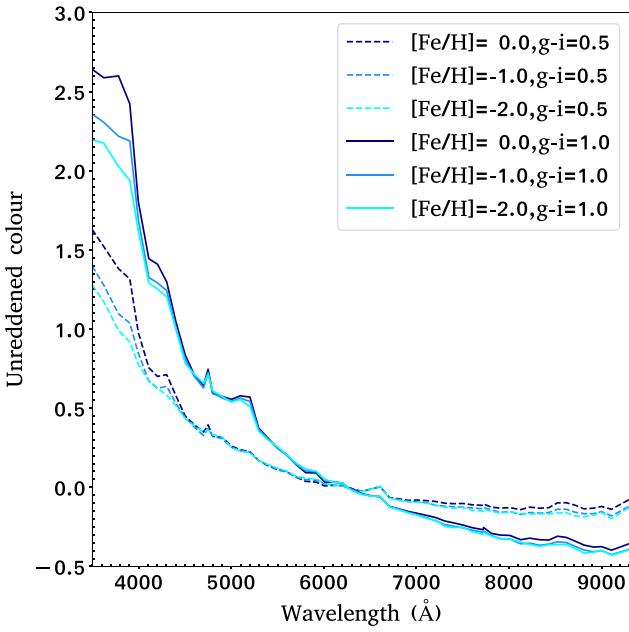


Figure 9. Comparison of $X - r$ colours between stars of the same $g - i$ colour (dashed lines: $g - i = 0.5$, solid lines: $g - i = 1.0$) but with different metallicities, where X refers to a given miniJPAS filter.

by Yang et al. (2022) are also shown, represented by red filled circles. This figure shows that VOSA T_{eff} follows a similar trend with the $(g - i)_0$ colour to the one found by Yang et al. (2022), but extending to a wider range of temperatures. This, again, validate our methodology to derive effective temperatures for thousands of sources from miniJPAS SEDs with VOSA.

4 METALLICITY-DEPENDENT STELLAR LOCI OF FGK DWARF STARS

We also used the Yang et al. (2022) J-PLUS DR1 stellar parameter estimations to determine the metallicity-dependent stellar loci of the miniJPAS colours for FGK dwarf stars. After requiring $r < 18$, $0.3 < (g - i)_0 < 1.6$, $4000 < T_{\text{eff}} < 7000$ K, $\log g > 4.0$, and $\text{flag} = 0$ in all the miniJPAS filters, 310 FGK dwarf stars of high-quality photometry and stellar parameters are selected. These stars, including 10 with $[\text{Fe}/\text{H}] < -1.5$, are used to determine the metallicity-dependent stellar loci of the miniJPAS colours.

Using the 310 miniJPAS-JPLUS FGK dwarf sample above, we have carried out a global two-dimensional polynomial fit to a series of 59 miniJPAS colours ($(X - r)_0$, where X refers to one of the 56 + 4 miniJPAS filters, except for r) as a function of the $(g - i)$ colour and metallicity $[\text{Fe}/\text{H}]$. A few stars with photometric errors larger than 0.05 mag are excluded from this step. Note that all colours referred to hereafter are dereddened intrinsic values, with reddening from Schlegel, Finkbeiner & Davis (1998) and reddening coefficients from the Schlafly et al. (2016) extinction curve.⁸ For colours $u\text{JAVA} - r$, $J0378 - r$, $J0390 - r$, $J0400 - r$, and $u\text{JPAS} - r$, a third-order polynomial with 10 free parameters is adopted. For the remaining colours that are relatively insensitive to metallicity, a second-order polynomial of six free parameters is used to avoid overfitting. A two-sigma clipping is performed during the fitting process.

⁸The numbers are stored in the miniJPAS.Filter table, available online.

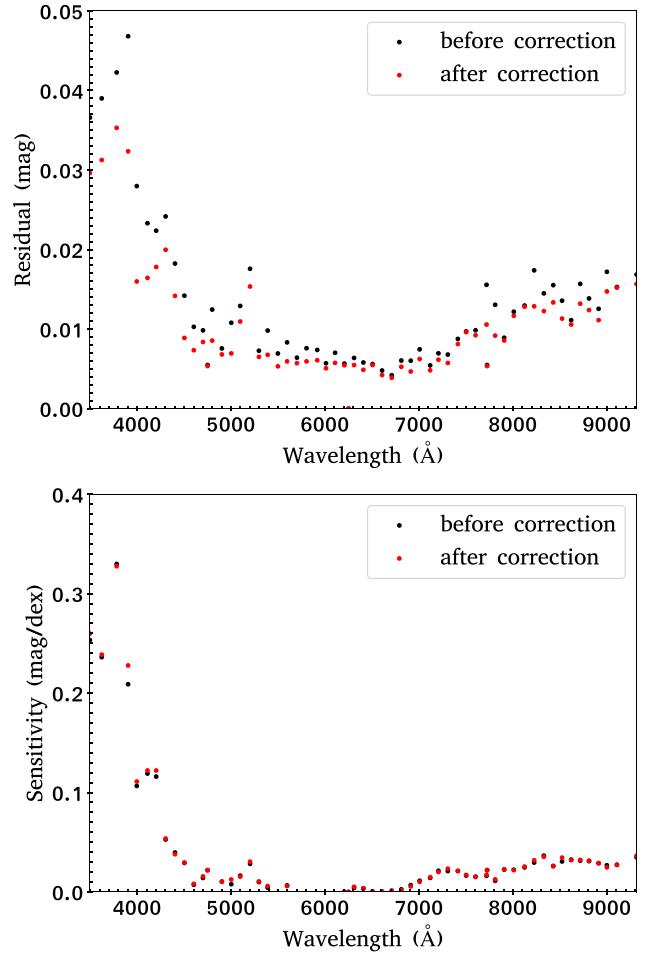


Figure 10. Top panel: Fitting residuals of metallicity-dependent stellar loci for different $X - r$ colours. Bottom panel: Typical metallicity sensitivities for different $X - r$ colours. The black and red dots represent results before and after the correction, respectively.

The fitting residuals show clear spatial variations in the CCD focal plane, from about 0.2 per cent in the red colours to a few per cent in the blue colours. These spatial variations are probably caused by flat-fielding (illumination correction) errors in the miniJPAS data. Therefore, we have performed a 2D polynomial fitting as a function of CCD (X, Y) positions to the initial locus residuals. The fitting coefficients are listed in Table 1. The results of flat-fielding errors are displayed in Fig. 6. The fitting residuals after correction are displayed in Fig. 7. In the next step, we have revised the miniJPAS colours/magnitudes accordingly to correct for their flat-fielding errors, assuming that there are no errors in the r band.⁹ The metallicity-dependent stellar loci of different miniJPAS filters are fitted again. The results are shown in Fig. 8. The resultant fitting coefficients are listed in Tables 2 and 3. The differences between the initial fitting curves and those after one iteration are, on average, smaller than 0.01 mag, so no further iterations are needed. Fig. 9 compares colours between stars of the same $g - i$ colour but with different metallicities, which are estimated using the fitted polynomials. The large differences in the blue filters are seen.

⁹Note that the revised magnitudes have also been used in the SED fittings with VOSA in Section 3.

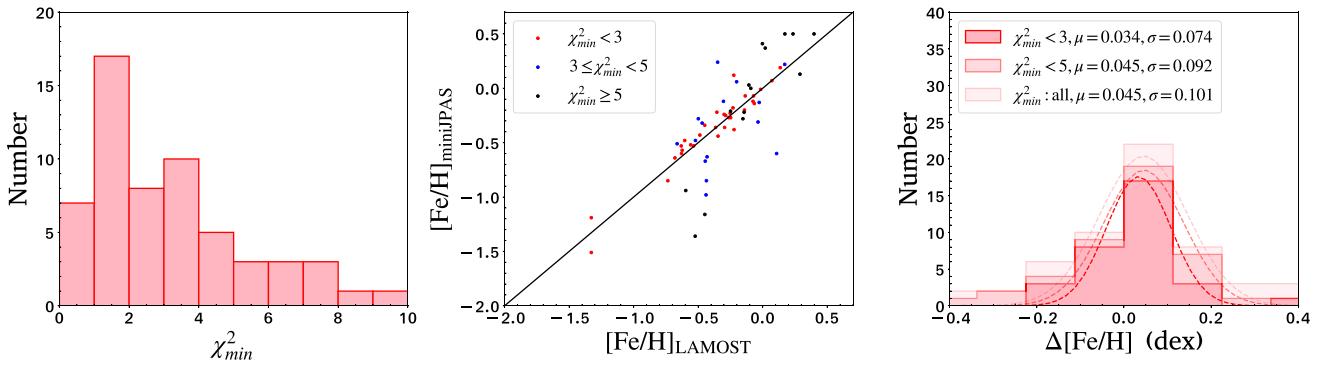


Figure 11. Left-hand panel: histogram distribution of the minimum χ^2 for the LAMOST-miniJPAS test sample. Middle panel: comparison between miniJPAS and LAMOST metallicities. The black line is the identify function. Right-hand panel: histogram distributions of the metallicity differences. The Gaussian fitting results are overplotted, with the mean and sigma values labelled.

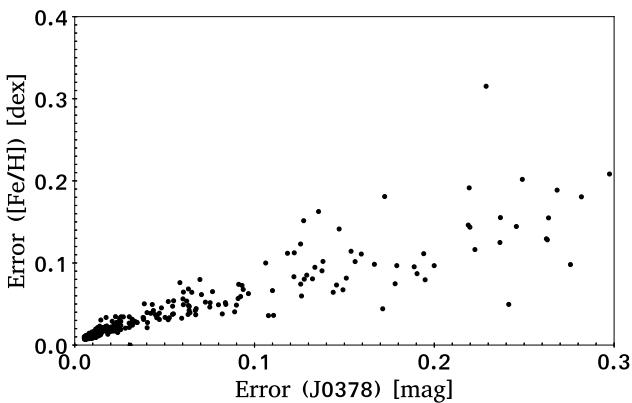


Figure 12. [Fe/H] errors as a function of $J0378$ errors for stars of $[Fe/H] > -1$.

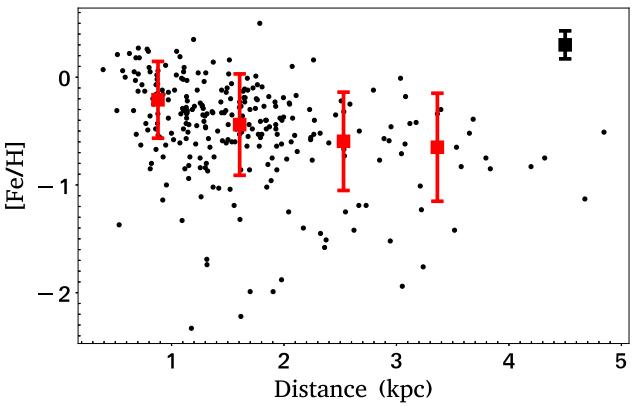


Figure 13. [Fe/H] as a function of distance for 269 selected dwarfs. See more details in Section 5. The red dots with error bars denote the median values and standard deviations of the black dots. The black error bar denotes the typical error of $\sigma_{[Fe/H]} = 0.13$ dex.

More details are plotted in Fig. 10. The top panel of Fig. 10 compares the locus fitting residuals before and after correction. It can be seen that the residuals decrease remarkably after correction, particularly for the blue colours. The fitting residuals are larger in the bluer colours for two main reasons: the stronger dependence on metallicity and the larger photometric errors. The larger numbers in

the redder colours are due to similar reasons. The typical residuals are around 0.005 mag, suggesting that the photometric quality of the miniJPAS data is very high. Note that there is a strong peak around the $J0510$ filter, which is caused by the presence of magnesium absorption features. Due to the Mg II triplet, the $J0510$ and $J0520$ filters show strong dependence on the Mg abundance (see Yang et al. 2022 for more details). Therefore, variations of [Mg/Fe] for a given [Fe/H] cause the larger fitting residuals. For similar reason, there is another peak in the $J0430$ filter, caused by the effect of [C/Fe] abundance and CH G-band feature. The peak at the $J0390$ filter is likely caused by the [Ca/Fe] abundance and CaII HK lines. The peak at the $J0378$ filter is probably due to its strongest metallicity dependence. The peaks at the $J0390$, $J0430$, $J0510$, and $J0520$ filters suggest that we can estimate individual element abundances such as [Ca/Fe], [C/Fe], and [Mg/Fe] from the miniJPAS/JPAS data, as already demonstrated with the JPLUS data by Yang et al. (2022).

The bottom panel of Fig. 10 shows the typical metallicity sensitivities for different colours, which are defined as the median variation of a given colour loci of $[Fe/H] = -1$ and $[Fe/H] = 0$. The very blue colours, including $uJAVA - r$, $J0378 - r$, $J0390 - r$, $uJPAS - r$, show very strong metallicity dependence, around 0.25 mag dex^{-1} . The sensitivities decrease to about 0.1 mag dex^{-1} for the $J400 - r$, $J410 - r$, and $J420 - r$ colours, and then to 0.01 mag dex^{-1} . The numbers increase to 0.03 mag dex^{-1} for colours redder than 7000 \AA .

5 PHOTOMETRIC METALLICITIES OF FGK DWARF STARS FROM MINIJPAS

In this section, we combine the empirical metallicity-dependent stellar loci of Section 4 and a minimum χ^2 technique to determine photometric metallicities of FGK dwarf stars from the miniJPAS data. Similar approaches have been applied to the SDSS/Stripe 82 (Yuan et al. 2015b; Zhang et al. 2021), *Gaia* EDR3 (Xu et al. 2022), and SMSS DR2 (Huang et al. 2022). Here, the χ^2 is defined as

$$\chi^2([Fe/H], g - i) = \sum_{i=1}^{59} \frac{[c_{\text{obs}}^i - R_c^i \times E(B - V) - c_{\text{int}}^i([Fe/H], g - i)]^2}{(\sigma_c^i)^2 \times (59 - 2)}, \quad (1)$$

where c_{obs}^i ($i = 1-59$) are the observed miniJPAS colours. R_c^i , $c_{\text{int}}^i([Fe/H], g - i)$, and σ_c^i ($i = 1-59$) are, respectively, the reddening coefficients, the intrinsic colours predicted by the metallicity-dependent stellar loci for a given set of metallicity [Fe/H] and intrinsic colour $g - i$, and the uncertainties of observed colours. Values of

$E(B - V)$ are from the extinction map of Schlegel et al. (1998). Values of σ_i are estimated from the magnitude errors and calibration uncertainties. The calibration uncertainties are adopted as 0.005 mag for all colours. The number is chosen for two reasons: (1) it is the typical precision we can achieve for blue filters with our method (e.g. Huang & Yuan 2022); (2) it is very close to the minimum value of the fitting residuals after correction in Fig. 10. The calibration uncertainties are likely overestimated for filters around the r band, but it will not affect the main result of this work, as the calibration uncertainties are usually much smaller than the magnitude errors.

We use a brute-force algorithm to determine the optimal $[\text{Fe}/\text{H}]$ for each sample star. For a given sample star, the value of $[\text{Fe}/\text{H}]$ is varied from -2.5 to 0.5 at a step of 0.01 dex. A 1D array of 301 χ^2 values is calculated to find the minimum value of χ^2 , χ^2_{\min} , and the corresponding metallicity $[\text{Fe}/\text{H}]$. A Monte Carlo method is used to estimate the errors of $[\text{Fe}/\text{H}]$ by introducing a Gaussian random noise for each colour 200 times.

We apply the method to 683 FGK stars of high quality photometry ($0.3 < g - i < 1.6$, $\text{flag} = 0$ in all filters, $rmag < 20$), assuming that they are all dwarfs.¹⁰ The median of the χ^2_{\min} values is 1.4. The typical χ^2_{\min} values are slightly larger when the stars are brighter, suggesting that the uncertainties are slightly underestimated for bright stars. To test the precision of the method, we have selected 61 dwarf stars in common with LAMOST DR7 and having LAMOST metallicity uncertainties smaller than 0.15 dex. The left-hand panel of Fig. 11 shows the histogram distribution of the χ^2_{\min} values for the LAMOST-miniJPAS test sample. The median value is about 2.6. The middle and right-hand panels compare the miniJPAS and LAMOST metallicities. A good agreement is seen, particularly for those of $\chi^2_{\min} < 3$. The metallicity differences have a mean value close to zero and a sigma value of about 0.1 dex, suggesting that the errors of our photometric metallicities are smaller than 0.1 dex. The results are not surprising, given the strong dependence on metallicity of the blue miniJPAS filters and the high-quality of the data. Note in the middle panel of Fig. 11, there are several stars having miniJPAS metallicities lower than LAMOST ones by 0.3 dex and $\chi^2_{\min} > 3$. They are probably binaries, as metallicities could be underestimated significantly for certain binaries by our method (e.g. Yuan et al. 2015b; Xu et al. 2022). One would expect that the $[\text{Fe}/\text{H}]$ errors depend on photometric errors, stellar colour, and $[\text{Fe}/\text{H}]$. Fig. 12 plots $[\text{Fe}/\text{H}]$ errors estimated from Monte Carlo simulations as a function of errors in the $J0378$ band for stars of $[\text{Fe}/\text{H}] > -1$. A tight linear relation is found. Similar relations are found for other filters, but with steeper slopes for filters that are less sensitive to $[\text{Fe}/\text{H}]$. The results demonstrate the power and potential of miniJPAS/JPAS in precise metallicity estimates for a large number of stars.

We cross-match the 683 stars with the *Gaia* EDR3 and further select dwarf stars ($M_r > 10*(g - i) - 4.0$ or $M_r > 4.0$) of $\chi^2_{\min} < 5$, $\text{PARALLAX} > 0$, and $\text{PARALLAX_OVER_ERROR} > 5$, where M_r is absolute magnitude in the r band. Fig. 13 plots metallicities of the selected 269 dwarf stars as a function of distance. A clear metallicity gradient is found, as expected. Note that stars of $[\text{Fe}/\text{H}] < -1.5$ may suffer large uncertainties due to the loci of < -1.5 are not well-constrained yet. There are a few stars with distances less than

¹⁰Note that we expect a few per cent giant stars in the sample. In this work, we did not obtain metallicity-dependent stellar loci for giant stars due to the lack of sufficient sample. Stellar loci are usually different between dwarfs and giants (e.g. Zhang et al. 2021). With up-coming JPAS data, the differences can be used to discriminate giants from dwarfs, as we did to the SDSS/Stripe 82 (Zhang et al. 2021).

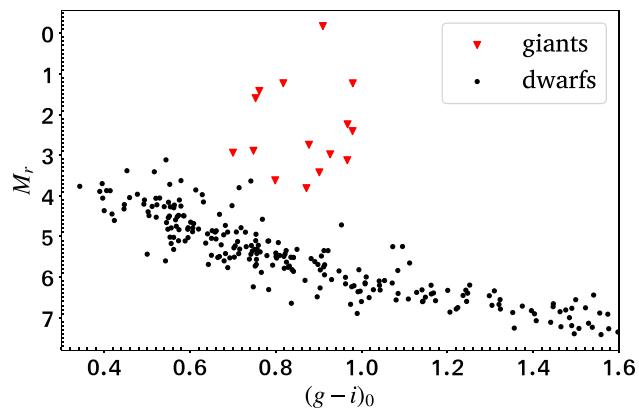


Figure 14. Distribution of the giant stars (red triangles) and dwarf stars (black dots) in the HR diagram.

1.5 kpc and $[\text{Fe}/\text{H}] < -1.5$. They are more likely binary stars in the disc rather than metal-poor stars in the halo, for reasons mentioned above that metallicities could be underestimated significantly for binaries. Besides, most of them have tangential velocities similar to those of disc stars (Xu et al. 2022).

6 GIANT DWARF CLASSIFICATION

Stellar loci between dwarf stars and giant stars are different, particularly in the blue filters (e.g. Zhang et al. 2021) and certain narrow-band filters (e.g. Morrison et al. 2000). However, due to the very limited giant stars available in the miniJPAS data, it is very challenging to determine metallicity-dependent stellar loci of giant stars. With the help of *Gaia* parallaxes, we are able to explore the potential of giant/dwarf classification with the 60 miniJPAS filters in this section.

The miniJPAS stars are cross-matched with the *Gaia* EDR3. In this work, we only use stars with high-precision photometry. The following criteria are required: (1) $\text{Flag} = 0$ in all miniJPAS filters; (2) $\text{ERR_APER6} < 0.25$ in all miniJPAS filters; and (3) $\text{PARALLAX_OVER_ERROR} > 9.0$ for dwarfs, > 1.0 for giants due to their much larger distances and consequently larger relative parallax errors.

Dwarf stars are selected empirically as those of $M_r > 10*(g - i) - 4.0$ or $M_r > 4.0$, while giant stars of $M_r < 10*(g - i) - 4.0$ and $M_r < 4.0$. Finally, we have obtained a sample of 233 dwarf stars and 15 giant stars. Their distribution in the HR diagram is shown in Fig. 14. Their distributions in different colour–colour diagrams are shown in Fig. 15.

The extreme gradient boosting (*XGBoost*; Chen & Guestrin 2016) is chosen as our classification model for the giant/dwarf stars. This algorithm can select important features during training process and mitigate overfitting by pruning, which improves performance of small sample classification. All the J-PAS filters are used. The training process is carried out with the PYTHON wrapper of XGBoost provided by scikit-learn. The best values of the hyperparameters are determined by the grid search with five-fold cross validation, and values ('learning_rate': 0.02, 'n_estimators': 90, 'max_depth': 4, 'min_child_weight': 1, 'seed': 0, 'subsample': 0.9, 'colsample_bytree': 0.6, 'gamma': 0.02, 'reg_alpha': 0.00009, 'reg_lambda': 0.0008) are used as the final hyperparameters in this paper.

The total samples are divided into the training set and the test set according to the ratio of 3:1. Both the training and testing sets are

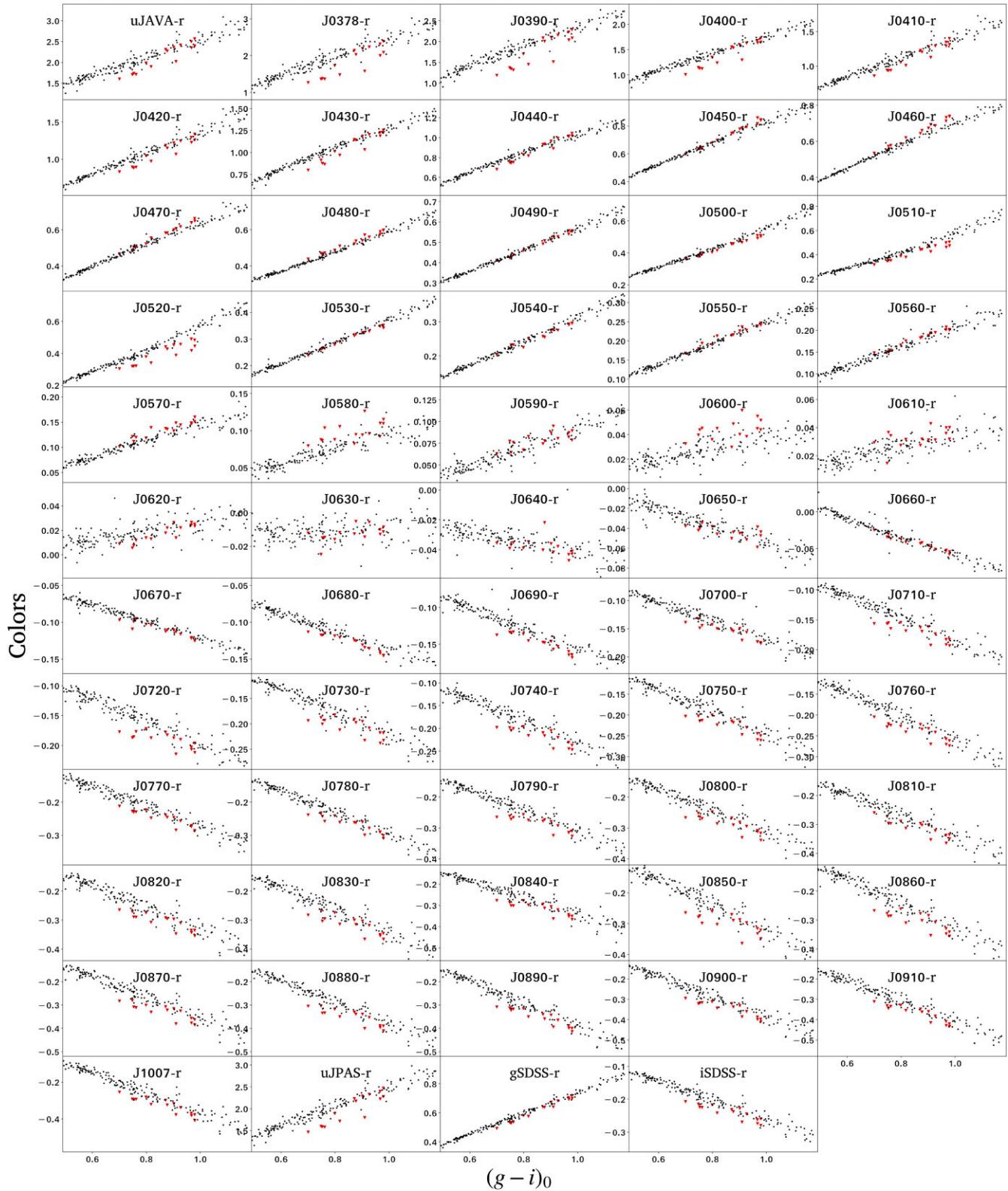


Figure 15. Distributions of the giant stars (red triangles) and dwarf stars (black dots) in the colour–colour diagrams. Note that the ranges in y-axes are different.

classified correctly, suggesting the strong potential of the miniJPAS filters in discriminating dwarfs and giants. We rank all the features (shown in Fig. 16) according to the total gain when the algorithm splits candidates. It can be seen that J0520 and J0510 are the two

most important filters, as expected, because the Mg triplet is more prominent in dwarfs than in giants for FGK stars. At the same $g - i$ colour, giants are usually bluer in $J0510 - r$ and $J0520 - r$ colours. Note that Mg abundances also show moderate effect on $J0510 - r$

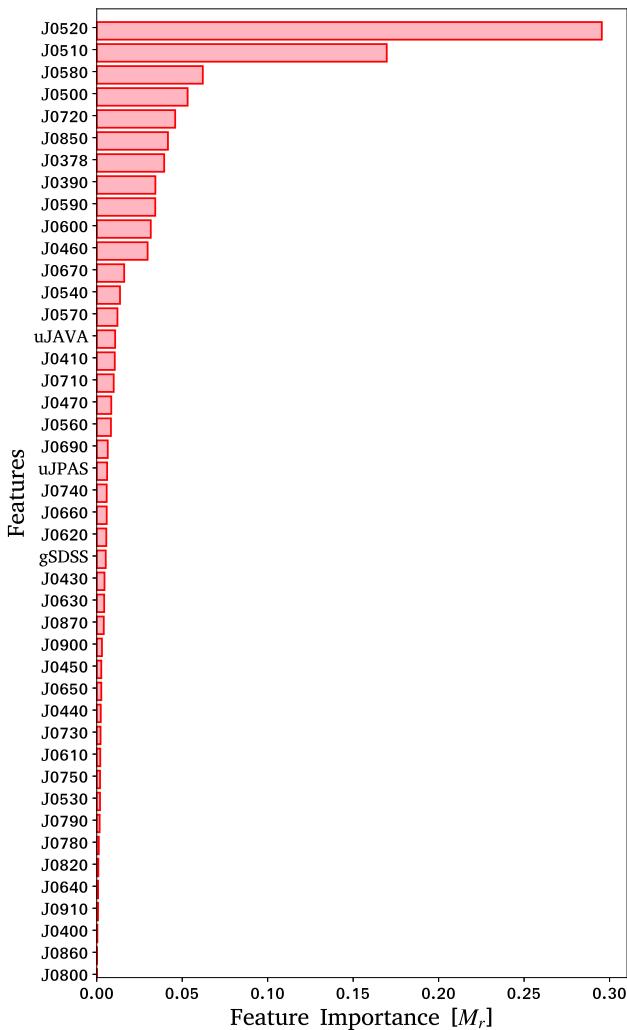


Figure 16. Relative importance of the features for the giant/dwarf classification.

and $J0520 - r$ colours (Yang et al. 2022). Note also that due to the limited number and parameter space (e.g. in $[Fe/H]$) of both giants and dwarfs in the above classification, the ranking of the features is probably biased. Further studies are needed to address this problem when the J-PAS data are available.

7 CONCLUSIONS

The unique set of 54 overlapping narrow band and two broad filters of the J-PAS survey provides a great opportunity for stellar physics and Galactic studies. In this work, we use the miniJPAS data to explore and prove the potential of the J-PAS filter system in characterizing stars and deriving their fundamental parameters via different techniques. The main conclusions are as follows.

(i) We used VOSA to derive stellar effective temperatures via SED fitting using miniJPAS photometry. We found that effective temperatures are in good agreement with those obtained from spectroscopic data from LAMOST and with previous works. We estimated a typical uncertainties < 150 K, which validate the high quality of the J-PAS data and postulates SED fitting technique as a

suitable way to obtain reasonably accurate effective temperatures for the million of stars expected to be observed in the J-PAS survey.

(ii) We have constructed the metallicity-dependent stellar loci in 59 colours for the miniJPAS FGK dwarf stars. The locus fitting residuals show significant spatial variations in the CCD focal plane, from about 0.2 per cent in the red colours to a few per cent in the blue colours. Such patterns, probably due to errors in the flat-fielding process, have been corrected. The very blue colours, including $uJAVA - r$, $J0378 - r$, $J0390 - r$, $uJPAS - r$, show the strongest metallicity dependence, around $0.25 \text{ mag dex}^{-1}$. The sensitivities decrease to about 0.1 mag dex^{-1} for the $J0400 - r$, $J0410 - r$, and $J0420 - r$ colours. The fitting residuals show peaks at the $J0390$, $J0430$, $J0510$, and $J0520$ filters, suggesting that we can estimate individual elemental abundances such as $[\text{Ca/Fe}]$, $[\text{C/Fe}]$, and $[\text{Mg/Fe}]$ from the miniJPAS/J-PAS data, as already demonstrated with the JPLUS data by Yang et al. (2022).

(iii) Combining the empirical metallicity-dependent stellar loci and a straight-forward minimum χ^2 technique, we have achieved a metallicity precision better than 0.1 dex solely from the miniJPAS photometry. The results demonstrate the power and potential of miniJPAS/J-PAS in precise metallicity estimates for a huge number of stars in the future. Given the strong dependence on metallicity of the blue miniJPAS/J-PAS filters and the high data quality, we expect to achieve a high precision for stars with a metallicity down to ~ -3.5 or lower in the future.

(iv) We have used *XGBoost* to classify dwarfs and giants. Both the training and testing samples are all successfully classified, suggesting the strong potential of the miniJPAS/J-PAS filters in discriminating dwarfs and giants. The $J0520$ and $J0510$ filters play an import role in the classifications, as expected.

The results obtained in this work demonstrate the power and potential of miniJPAS/J-PAS data in stellar parameter determinations. Due to the very small number of stars in the small miniJPAS footprint, scientific investigations are limited. Note a very nice exploration of white dwarf science using the miniJPAS data by López-Sanjuán et al. (2022). In the future with the J-PAS data, stellar parameters can be better determined, not only for effective temperature, metallicity, and surface gravity, but also for reddening, distance, age, $[\text{C/Fe}]$, $[\text{N/Fe}]$, $[\text{Mg/Fe}]$, and $[\text{Ca/Fe}]$. Such a magnitude-limited sample of stars down to $r \sim 21-22$ and within 8500 deg^2 will be valuable for studies of stellar populations, structure, chemistry, and evolution of the Galaxy.

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DATA AVAILABILITY

The data underlying this article is available online via http://www.j-pas.org/datarereleases/minijpas_public_data_release_pdr201912 and <http://www.j-plus.es/ancillarydata/index>. The VOSA temperatures as well as the metallicities in Section 5 are available from the authors upon request.

REFERENCES

Allard F., Homeier D., Freytag B., 2011, in Johns-Krull C. M., Browning M. K., West A. A., eds, *ASP Conference Series*, Vol. 448, 16th Cambridge Workshop on Cool Stars, Stellar Systems, and the Sun. Astron. Soc. Pac., San Francisco, p. 91

Bayo A., Rodrigo C., Barrado Y. Navascués D., Solano E., Gutiérrez R., Morales-Calderón M., Allard F., 2008, *A&A*, 492, 277

Benítez N. et al., 2014, preprint ([arXiv:1403.5237](https://arxiv.org/abs/1403.5237))

Bertin E., Arnouts S., 1996, *A&AS*, 117, 393

Bonoli S. et al. 2021, *A&A*, 653, A31

Cenarro A. J. et al., 2014, *Proc. SPIE Conf. Ser. Vol. 9149, Observatory Operations: Strategies, Processes, and Systems V*. SPIE, Bellingham, p. 91491I

Chen T., Guestrin C., 2016, preprint ([arXiv:1603.02754](https://arxiv.org/abs/1603.02754))

Gaia Collaboration et al., 2016, *A&A*, 595, A1

Gaia Collaboration et al., 2018, *A&A*, 616, A1

Gaia Collaboration et al., 2021, *A&A*, 649, A1

Huang B., Yuan H., 2022, *ApJS*, 259, 26

Huang Y. et al., 2021, *ApJ*, 907, 68

Huang Y. et al., 2022, *ApJ*, 925, 164

López-Sanjuán C. et al., 2022, *A&A*, 665, 13

Luo A.-L. et al., 2015, *Res. Astron. Astrophys.*, 15, 1095

Marín-Franch A. et al., 2012, *Proc. SPIE Conf. Ser. Vol. 8450, Design of the J-PAS and J-PLUS filter systems*. SPIE, Bellingham, p. 84503S

Marín-Franch A. et al., 2017, in Arribas S., Alonso-Herrero A., Figueiras F., Hernández-Monteagudo C., Sánchez-Lavega A., Pérez-Hoyos S., eds, *Proceedings of the XII Scientific Meeting of the Spanish Astronomical Society, Highlights on Spanish Astrophysics IX*. p. 670

Morrison H. L., Mateo M., Olszewski E. W., Harding P., Dohm-Palmer R. C., Freeman K. C., Norris J. E., Morita M., 2000, *AJ*, 119, 2254

Niu Z., Yuan H., Liu J., 2021, *ApJ*, 909, 48

Niu Z., Yuan H., Liu J., 2021, *ApJ*, 908, L14

Rockosi C. M. et al., 2022, *ApJS*, 259, 60

Schlafly E. F. et al., 2016, *ApJ*, 821, 78

Schlegel D. J., Finkbeiner D. P., Davis M., 1998, *ApJ*, 500, 525

Taylor K. et al., 2014, *J. Astron. Instrum.*, 3, 1350010–685

Xiao K., Yuan H., 2022, *AJ*, 163, 185

Xu S., Yuan H., Niu Z., Yang L., Beers T. C., Huang Y., 2022, *ApJS*, 258, 44

Yang L., Yuan H., Zhang R., Niu Z., Huang Y., Duan F., Fang Y., 2021, *ApJ*, 908, L24

Yang L. et al., 2022, *A&A*, 659, A181

York D. G. et al., 2000, *AJ*, 120, 1579

Yuan H., Liu X., Xiang M., Huang Y., Chen B., 2015a, *ApJ*, 799, 134

Yuan H., Liu X., Xiang M., Huang Y., Chen B., 2015b, *ApJ*, 803, 13

Yuan H., Liu X., Xiang M., Huang Y., Zhang H., Chen B., 2015c, *ApJ*, 799, 133

Zhang R.-Y., Yuan H.-B., Liu X.-W., Xiang M.-S., Huang Y., Chen B.-Q., 2021, *Res. Astro. Astrophys.*, 21, 319

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